atomic absorption and emission of light

1889 Rydberg generalized to other H emission lines:

$$\therefore \frac{mv^2}{r} = \frac{ke^2Z}{r^2} \qquad L = mvr = n \hbar \qquad E = \frac{1}{2}mv^2 - \frac{ke^2Z}{r}$$

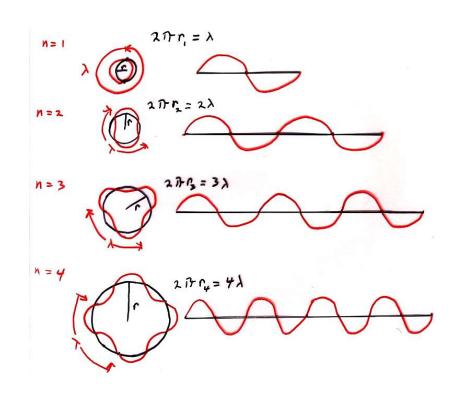
$$r = .53 \frac{n^2}{Z} (A)$$

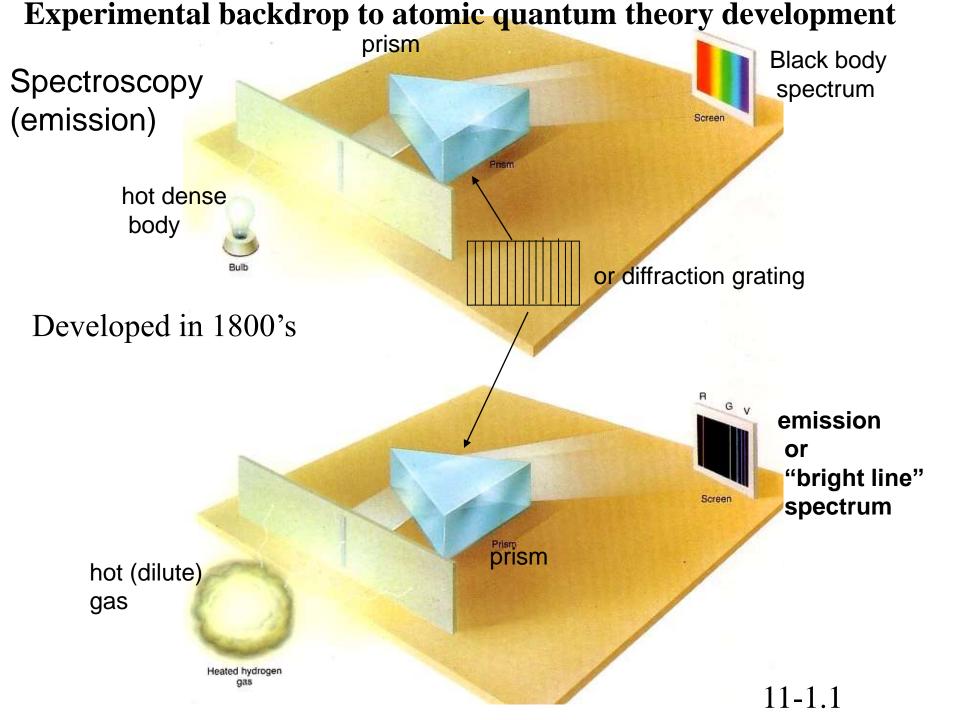
$$E_n = -13.6 \frac{(Z^*)^2}{n^2} (eV)$$

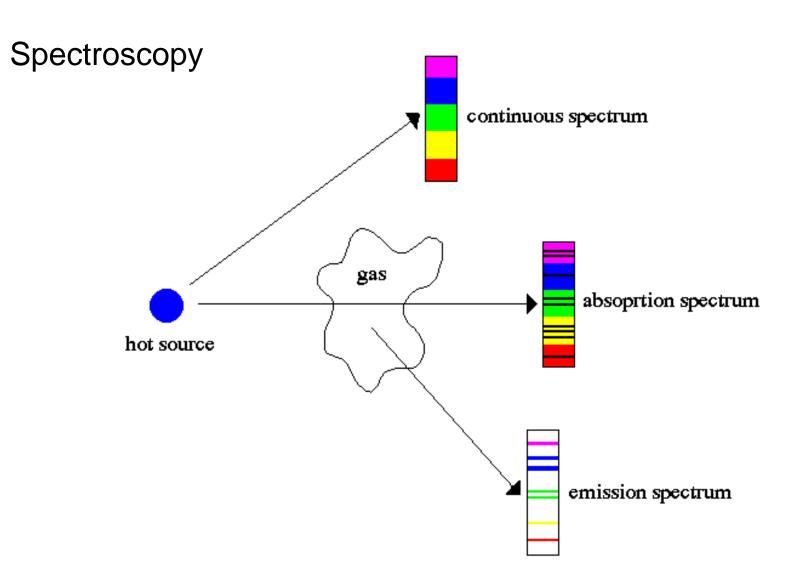
$$p = \frac{h}{\lambda}$$

Mosley's Law

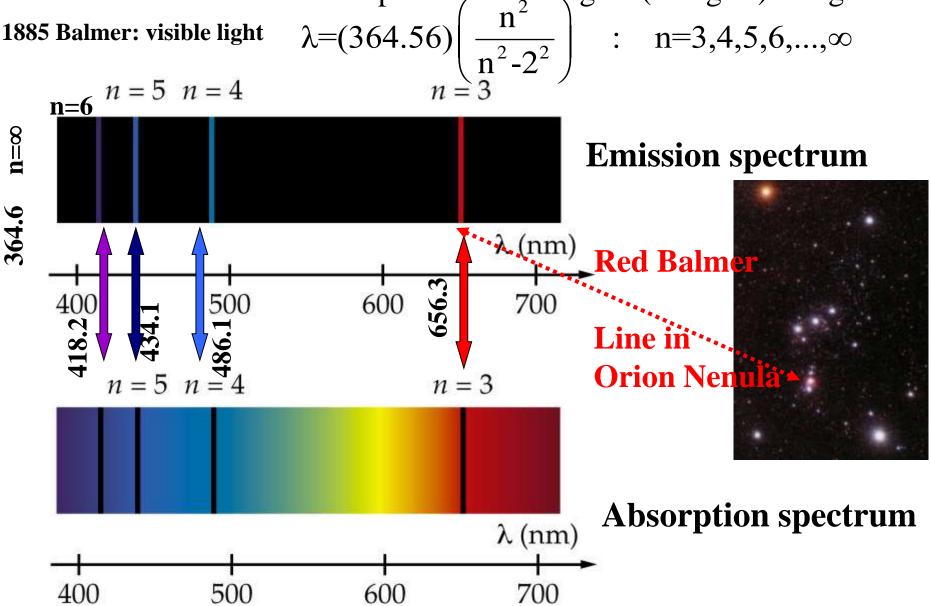
$$E_{f} = E_{n=2} - E_{n=1} *$$







H atom absorbs and emits specific wavelengths (energies) of light



1889 Rydberg generalized to other H emission lines:

Energy (discrete set) levels of H (all atoms) are <u>unique</u> ⇒ photons absorbed /emitted unique finger print of atom

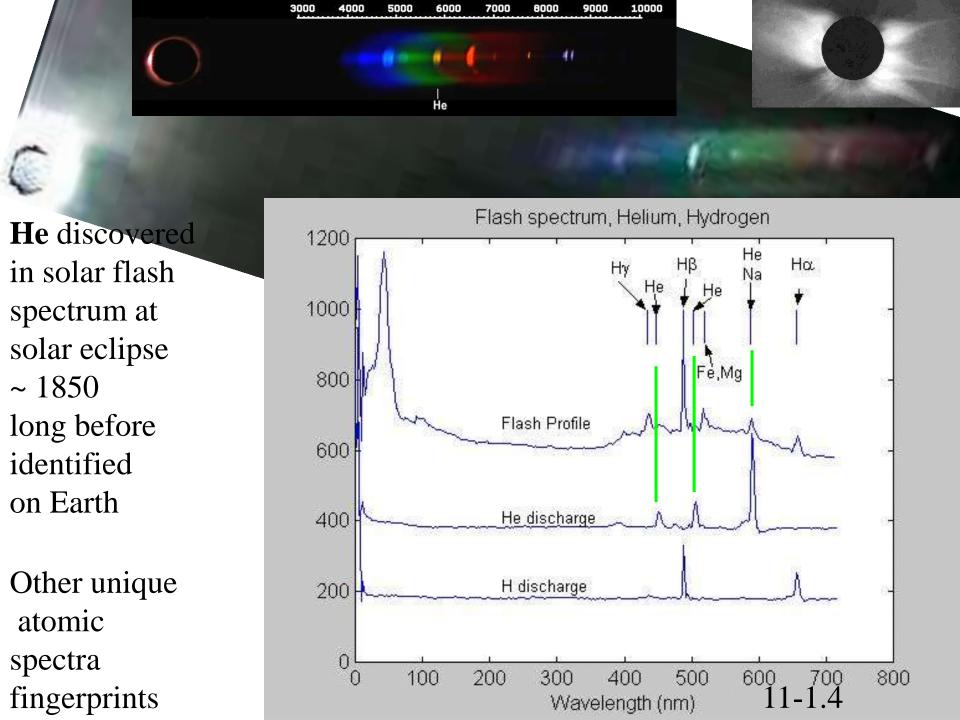
$$\frac{1}{\lambda} = \Re_H \left(\frac{1}{n_1^2} - \frac{1}{n_2^2} \right) : n_1 < n_2$$

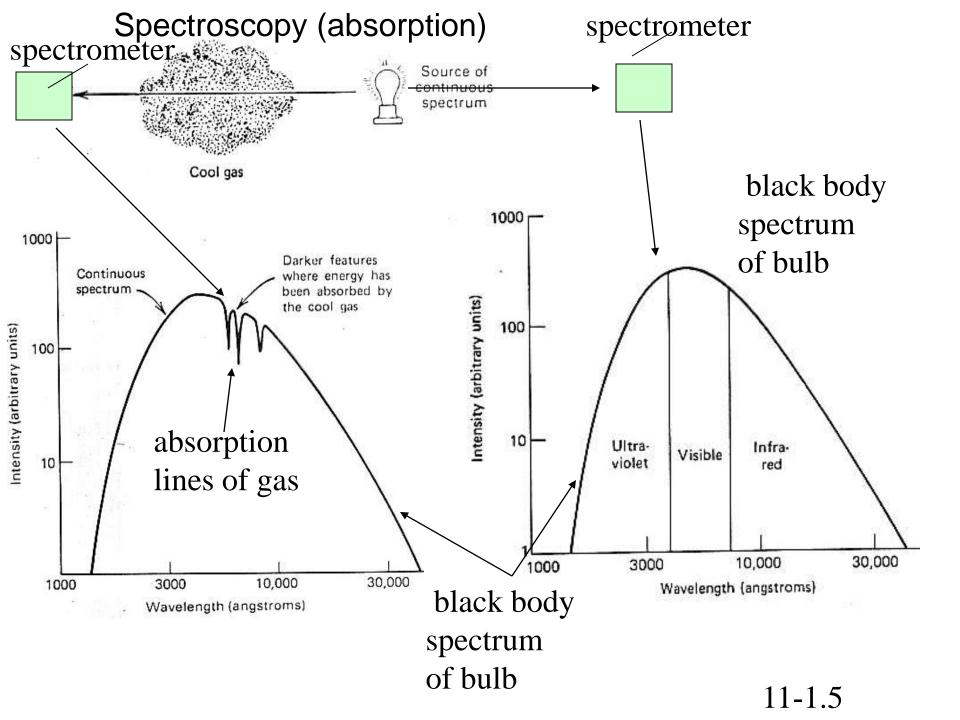
specifically
$$n_2 = n_1 + 1, n_1 + 2, \dots$$

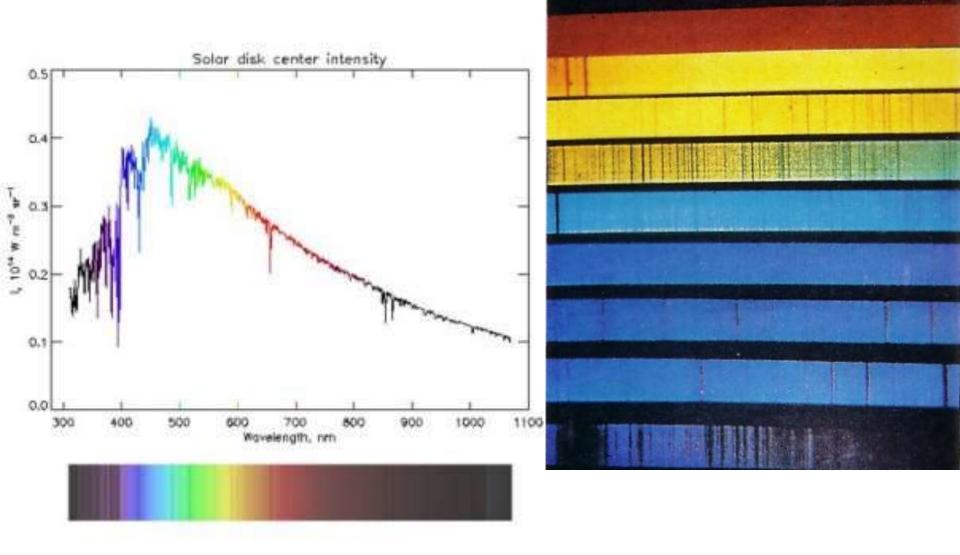
R_H is the *Rydberg constant*

$$(\mathfrak{R}_{H}=10,973,731.534 \text{ m}^{-1}).$$

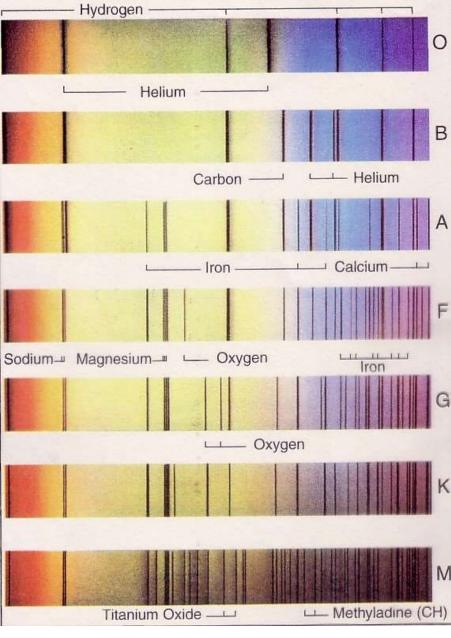
Energy discrete set of H emission lines are <u>unique</u> (and incredibly precise) finger print of H- atom







absorption lines in Sun's spectrum



absorption lines in spectra of stars with different surface temperatures (mass)

11-1.6a

Before 1913 – mountains of incredibly precise atomic spectral data. Atoms interact with light in highly organized discrete ways.

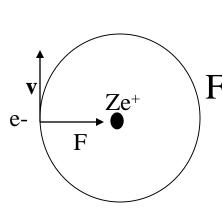
Only some wavelengths "fit".

Totally un-understandable in classical physics of particles.

1913 Bohr ("something stinks in Denmark")

Classical first step

Circular orbit



$$F = \frac{mv^2}{r}$$
 circular motion

$$F = \frac{\text{Inv}}{r} \text{ circular motion}$$

$$\frac{k(e)(Ze)}{r} \text{ Coulombs Law}$$

$$mv^2 = \frac{ke^2Z}{r}$$

1913 Bohr Atom

Circular

11-1.8a

can't change **Angular Momentum conserved** continuously (Bohr had to assume this!)

$$\begin{bmatrix} \frac{1}{2\pi} \end{bmatrix}$$

"stationary state"

Only quantized orbits allowed!!
$$r_{n} = n^{2} \left(\frac{\hbar^{2}}{\text{ke}^{2}\text{Zm}}\right) = \frac{n}{Z} [.53(10)^{-10}\text{m}] = \frac{1}{[.53\text{A}]} \cdot \frac{n^{2}}{Z} = r_{n}$$

J. Hand correction

2 Bohr radius

 \Rightarrow 1 A = atomic size

$$E=KE+PE=\frac{1}{2}mv^2-\frac{ke^2Z}{r}$$

$$\frac{|recall|}{mv^2 = \frac{ke^2Z}{r}}$$

$$E = \frac{1}{2} \left(\frac{Ke^{2}Z}{r} \right) - \frac{Ke^{2}Z}{r} = -\frac{1}{2} \frac{Ke^{2}Z}{r}$$

But recall
$$r_n = n^2 \left(\frac{\hbar^2}{ke^2 Zm}\right)$$

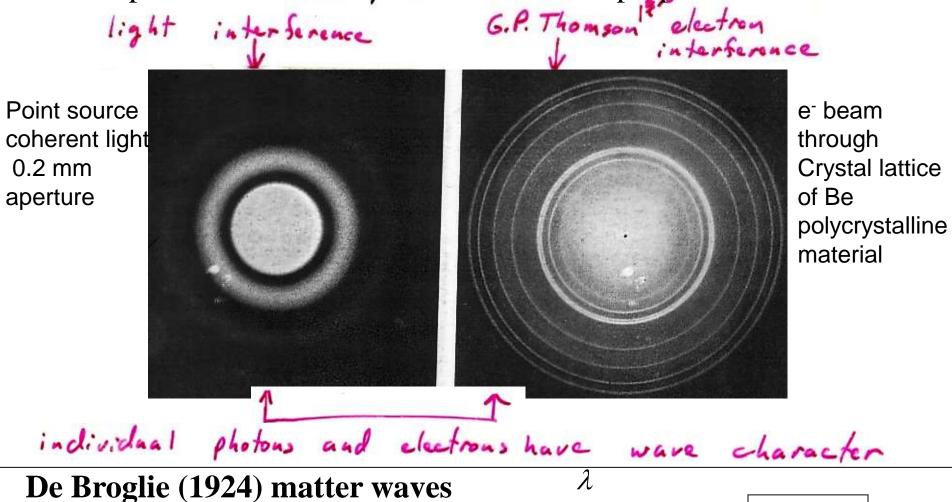
$$\therefore E = -\frac{1}{2} \frac{ke^2Z}{\frac{n^2\hbar^2}{ke^2Zm}}$$

$$E_n = -(\frac{k^2 e^4 Z^2 m}{2\hbar^2}) \frac{1}{n^2} = -(13.6 \text{ eV}) \frac{Z^2}{n^2}$$
 quantized orbits \Rightarrow quantized energy!

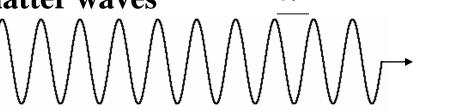
for Z=1 H
$$E_1 = -2.17(10)^{-18}$$
J=-13.6eV

Planck's law + Einstein photons= light has some "particle like" properties 1901 1905 PE effect

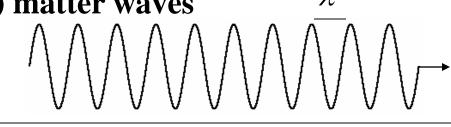
GP Thompson -1930 – electron has some wave properties



 $\begin{array}{ccc}
& & V \\
& & & \\
m & p = mv
\end{array}$



11-3



$$p = \frac{h}{\lambda}$$

Recall Bohr said for e⁻ in atomic orbit: 27r = x

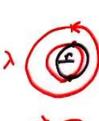
L= mvr = pr = n
$$\frac{h}{2\pi}$$

 \therefore if p= $\frac{h}{\lambda}$

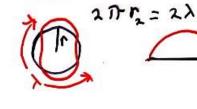
m p=mv

then
$$r \frac{h}{\lambda} = n \frac{h}{2\pi}$$

$$2\pi r = n\lambda$$

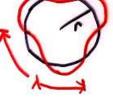


n = 1

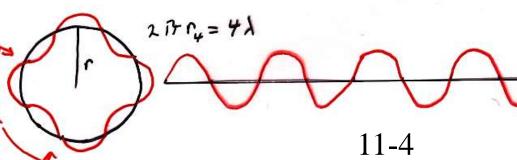












Bohr Theory – atomic emission spectra

$$E_n = -13.6 \frac{Z^2}{n^2}$$
 Z=atomic # Z=1 hydrogen
$$hf = E_{n2} - E_{n1}$$

$$\mathbf{n_2}$$

$$\mathbf{n_3}$$

Electron drops down
$$\Rightarrow$$
 photon emitted

$$\frac{13.6eV}{hc} = \frac{13.6 \cdot 1.6(10)^{-19} c \frac{J}{c}}{6.626(10)^{-34} Js \cdot 3(10)^{8} \frac{m}{c}} \qquad E_{n2} - E_{n1} = (13.6) \left[\frac{1}{n_{1}^{2}} - \frac{1}{n_{2}^{2}} \right] = fh$$

$$5.626(10)^{-34} Js \ \ 3(10)^{8} \frac{m}{s}$$

$$10^{19-8+34} m^{-1}$$

$$= \frac{(13.6)(1.6)}{(6.626)(3)} 10^{19-8+34} m^{-1}$$

$$= 10.973(10)^6 m^{-1} = R$$

$$(6.626)(3)$$

$$= 10.973(10)^{6} m^{-1} = R$$

$$\frac{1}{\lambda} = \{R\} \left[\frac{1}{n_{1}^{2}} - \frac{1}{n_{2}^{2}} \right]$$

$$c = \lambda f \qquad \therefore \qquad f = h \frac{c}{\lambda}$$

$$c = \lambda f \qquad \therefore \qquad f = h \frac{c}{\lambda}$$

$$h\frac{c}{\lambda} = (13.6)\left[\frac{1}{n_1^2} - \frac{1}{n_2^2}\right]$$

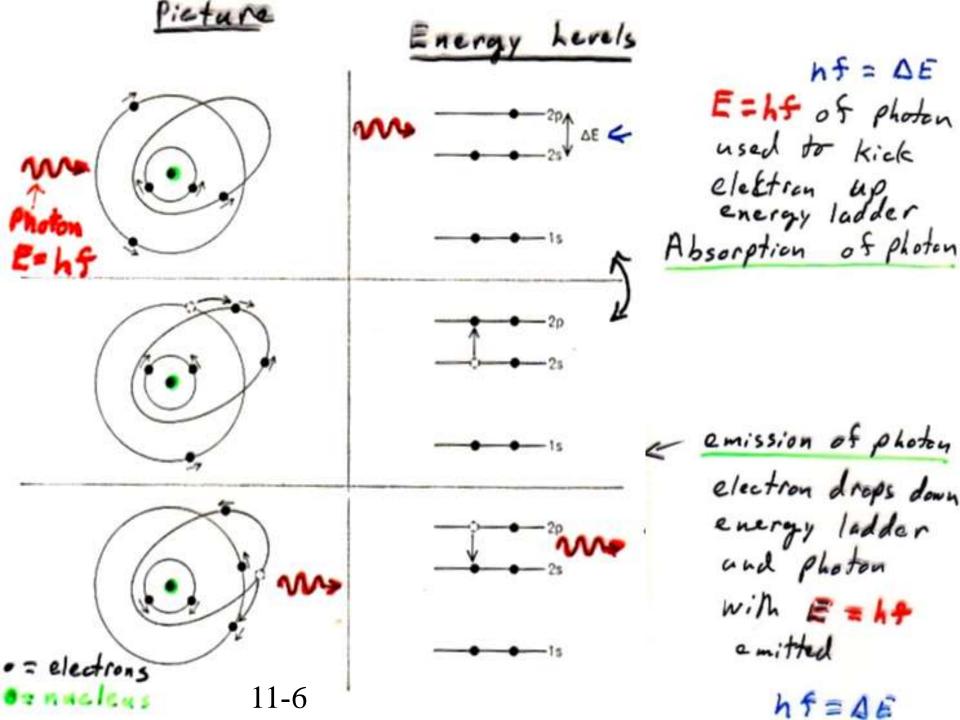
$$\frac{1}{\lambda} = \left\{\frac{(13.6)}{n_1^2}\right\}\left[\frac{1}{n_2^2} - \frac{1}{n_2^2}\right]$$

$$\frac{1}{\lambda} = \left\{ \frac{(13.6)}{hc} \right\} \left[\frac{1}{n_1^2} - \frac{1}{n_2^2} \right]$$

Rydberg formula from Bohr!!!

 $(\mathfrak{R}_{H} = 10,973,731.534 \text{ m}^{-1}).$ Precision of R unparalleled in history of physics !!!

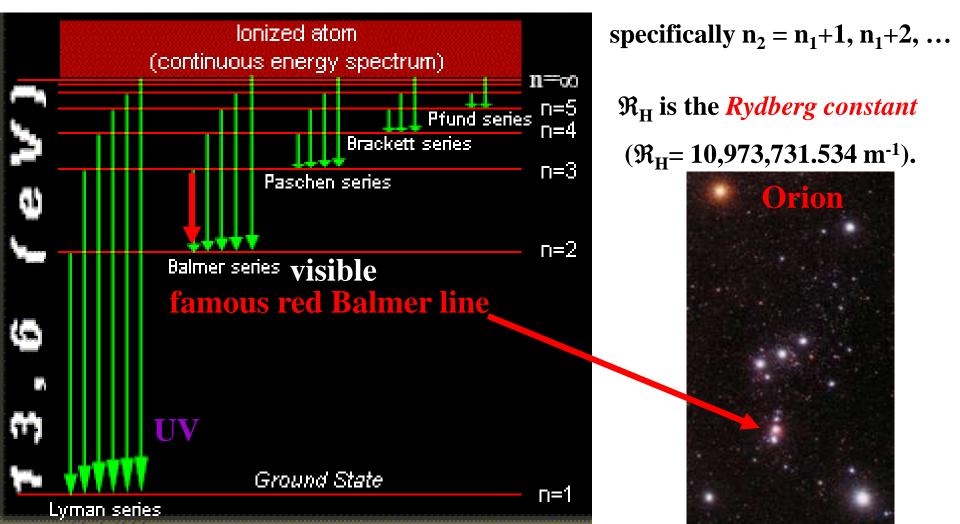
11-5

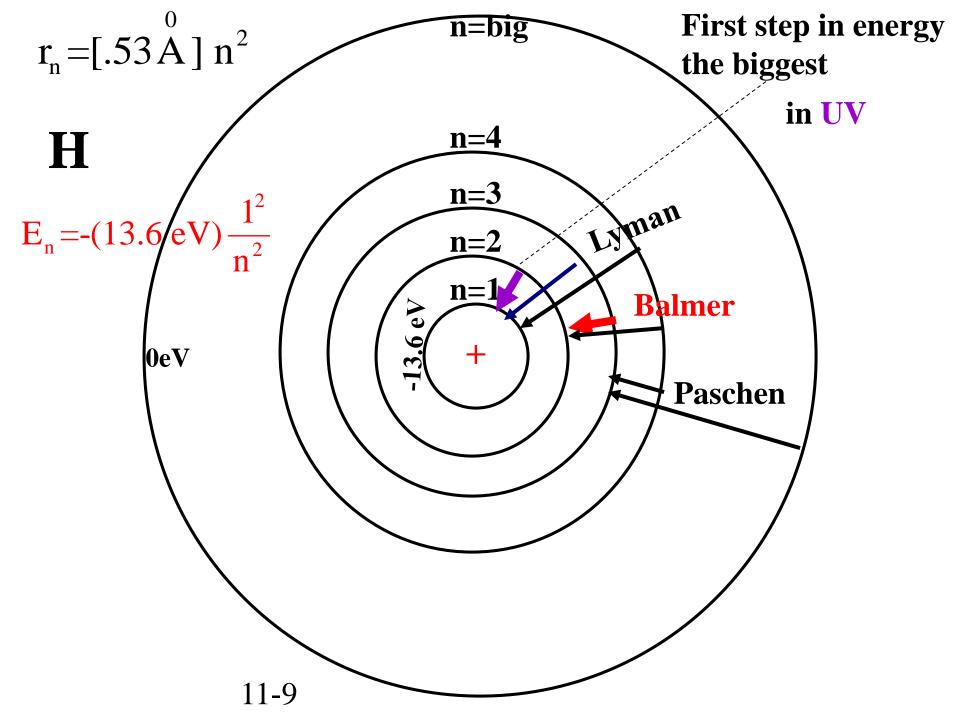


Ionization electron kicked out energy above ionization energy possible] electron drops

⇒ photons absorbed /emitted unique finger print of atom

1889 Rydberg generalized to other H emission lines: $\frac{1}{\lambda} = \Re_H \left(\frac{1}{n_1^2} - \frac{1}{n_2^2} \right) : n_1 < n_2$





X-Rays and X-Ray Emission

Roentgen 1895 X-rays (X-Strallung)

Max von Laue diffraction 1912-1913 (diffraction from planes of atoms in crystal)

Barkla 1911 x-ray emission – "hardness" = penetrating power hard & soft x-rays different targets – differing x-ray hardness emitted (gave us K, L, M, N x-ray emission lines)

Bragg 1912/13 x-ray diffraction $n\lambda=2dsin(\theta)$ [can use θ to measure interatomic distances d in solids !!!!]

Mosely (demonstrator in Rutherford's lab at Manchester)

- Bohr visits Rutherford Mosely told of new Atomic theory
- 1913- Mosely used Bragg reflection to determine λ of x-ray emission of various elements $n\lambda=2d\sin(\theta)$ --- know d of crystal measure θ to determine λ

1913- Mosely's Law
$$1/\lambda = K [Z-\sigma]^2$$
 $\sigma_K=1 \sigma_L=1$ Z=atomic number !!!!!

"Mosely and the Numbering of the Elements"- by Jaffe – Doubleday NY 1971

1915 - Mosely- killed in WWI – Dardanelles Campaign- (against Turkey)
{Winston Churchhill's idea (1'st Lord of the Admiralty)
– blood bath and WC didn't come back till WWII}

11-10 mos1

