

**(Astro)Physics 343 Lecture # 9:  
telescopes; interferometry**

# This week's schedule

**Monday – Thursday: ad hoc office hours for lab # 4  
(Baker for sections A, B, F; Naudus for C, D, E)**

**Monday + Thursday: regular office hours**

**Lab # 4 due next Monday.**

**Still waiting for responses from three people re trip.**

# Why do we need telescopes?

If a simple dipole antenna can detect radio waves...



long-wavelength  
development array  
at VLA site (test  
bed for larger array  
with  $\lambda=0.3-30\text{m}$ )

**Answer: collecting area (for short-wavelength observations).**

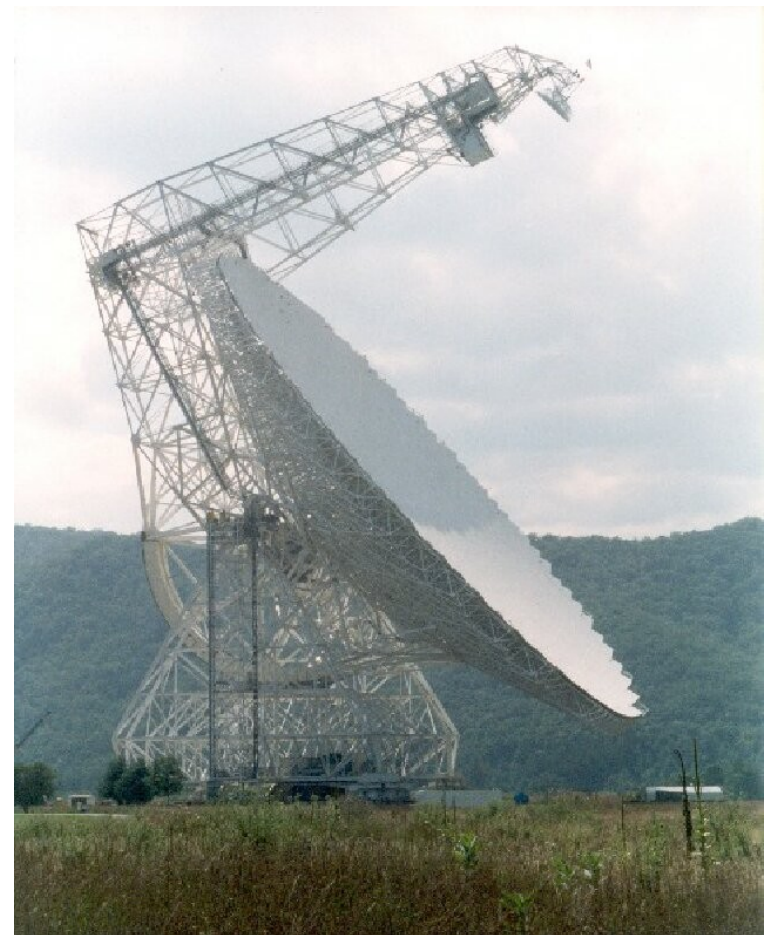
**Telescopes collect and focus power onto a smaller (e.g.,  
feed horn or dipole) antenna.**

# Telescope designs: feed horn vs. paraboloid

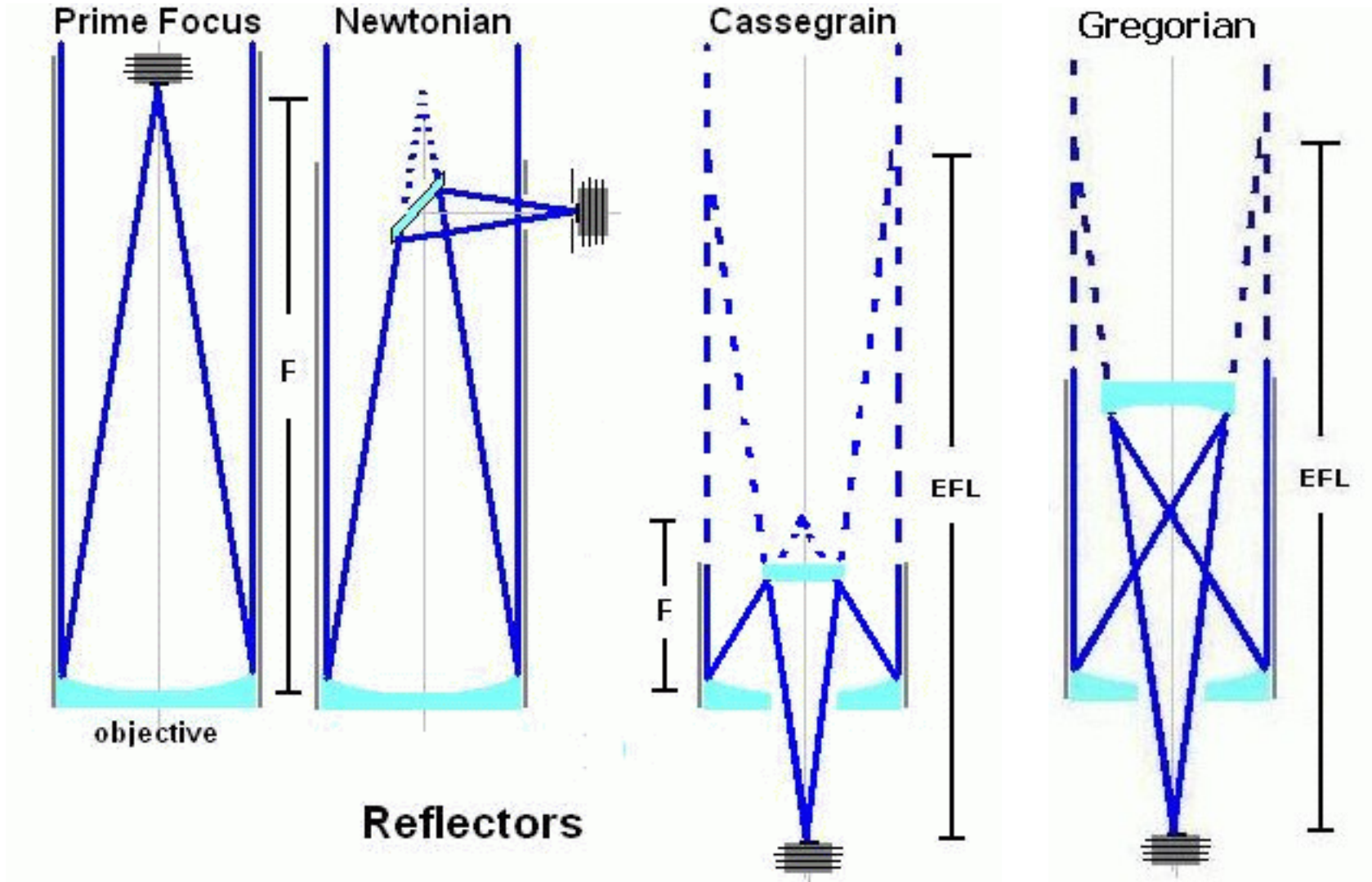


**feed horns: response can be calculated a priori! but size limited...**

**paraboloid antennas: good for collecting area, calibration tricky**



# Telescope designs: location of foci



**Borrowed from J. Oliver.**

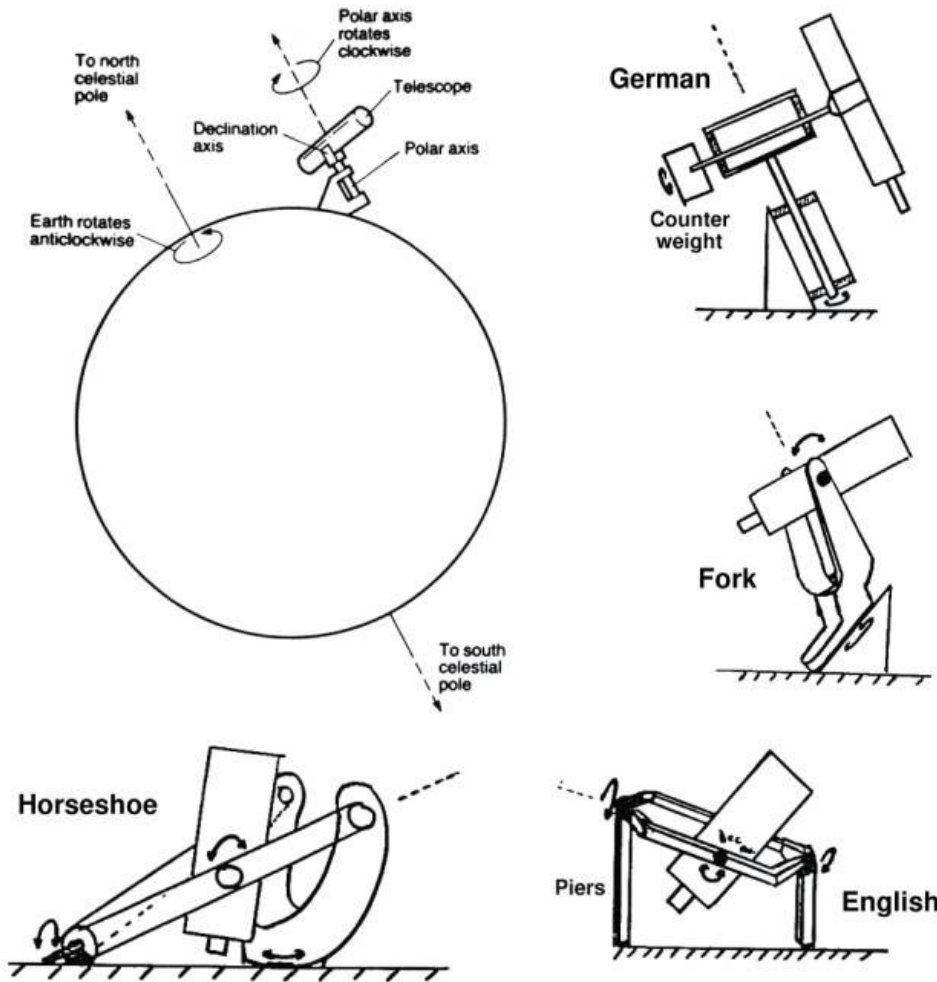
# Telescope designs: on or off axis?



**100m Effelsberg telescope (Germany) + Green Bank Telescope (WV)**

# Telescope designs: mount?

**alt-az: both axes to track sources**



**equatorial: one axis to track sources**

# The last big scope with an equatorial mount



**140 ft telescope at Green Bank:**

**(1) world's largest telescope  
with an equatorial mount**

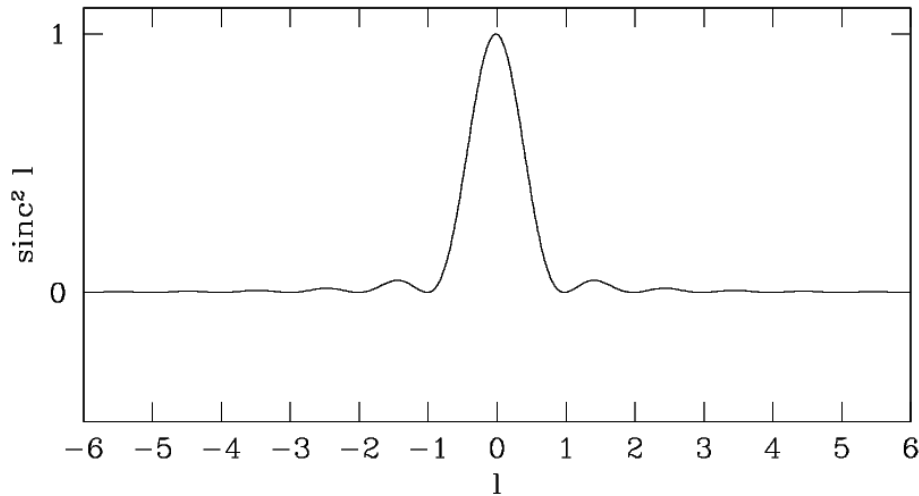
**(2) contains world's largest  
ball bearing!**



# The (angular) resolution of a telescope

We know that the FWHM of the telescope's beam is proportional to  $\lambda/D$ .

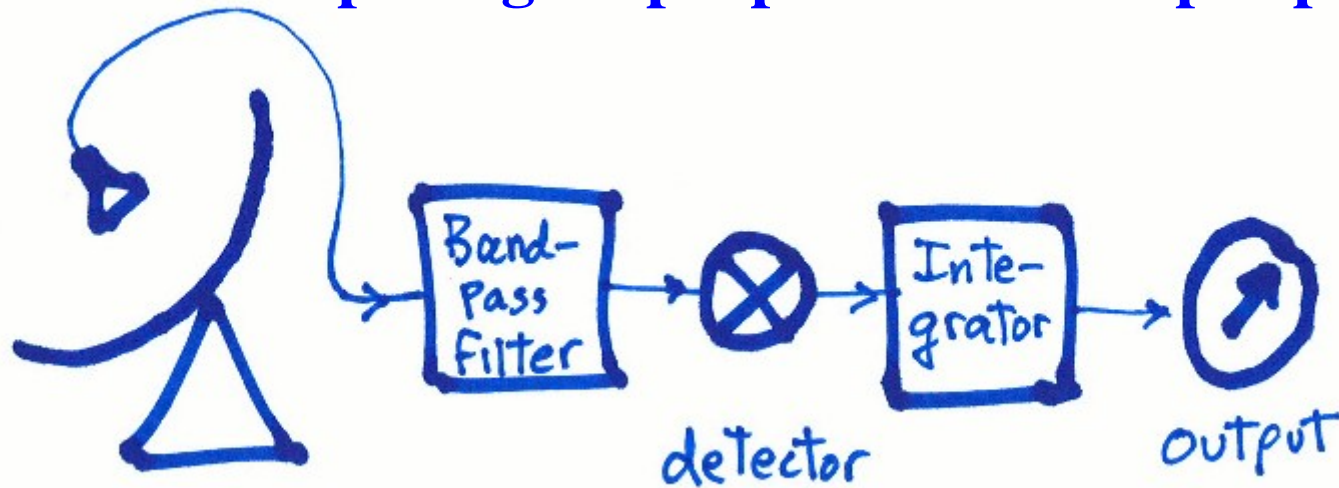
By the Rayleigh criterion, this is also its resolving power: two point sources separated by the FWHM will sit on peak + first dip of response.



# Radiometers

Steps in detection of radio emission with a **radiometer**:

- (1) select a frequency bandpass
- (2) multiply signal by itself
- (3) integrate over some time interval
- (4) record output signal proportional to input power



**Borrowed from Condon & Ransom, ERA.**

# Sensitivity of an ideal radiometer

If system temperature is  $T_{\text{sys}}$

... bandwidth is  $\Delta\nu$

... integration time is  $\Delta t$

then the sensitivity ( $1\sigma$  noise) will be

$$\Delta T = T_{\text{sys}} / \text{sqrt}(\Delta\nu \Delta t)$$

i.e., goes down as  $\text{sqrt}(\text{number of samples})!$

# Quiz